Table 1: Known hot cores and their physical parameters. Mass and luminosity estimates are derived respectively from (sub)millimeter continuum emission and from infrared flux densities.

Table 1

Source Name	Dist. (kpc)	$T_{\rm kin}$ (K)	Diam. (pc)	Mass ($M_\odot$)	$L_{\rm IR}$ ($L_\odot$)	Ref. code
Orion-KL	0.5	300	0.05	10	1.0 10^5	1
SgB2N[a]	8.5	200	0.1	2000	6.5 10^6	2 3
SgB2M[a]	8.5	200	0.1	2000	''	2
G5.89–0.39[a]	4.0	90	0.18	2800	7.1 10^5	4 5
G9.62+0.19	5.7	>100	0.059	55–160	4.4 10^5	6 7
G10.47+0.03[a]	5.8	150	0.078	2200	5.0 10^5	6 8 9 10
G10.62–0.38[a]	6.0	144	0.05	1100[b]	1.1 10^6	11 12
G19.62–0.23	3.5	230	0.032	450	1.6 10^5	13
G29.96–0.02	7.4	100	0.052	460	1.4 10^6	6 8 14
G31.41+0.31	7.9	110	0.080	2500	2.6 10^5	6 8 9
G34.26+0.15	3.8	250	0.066	1400	6.3 10^5	4 15 16 17 18
G45.07+0.13[a]	8.3	140	0.027	<7800	1.1 10^6	19
G45.12+0.13[a]	8.3	120	0.056	<37000	1.3 10^6	19 20
G45.47+0.05[a]	8.3	90	0.056–0.3	250	1.1 10^6	20
IRAS 20126+4104	1.7	200	0.012	10	1.3 10^4	21
DR 21(OH) MM1	3.0	>80	0.05	350	5.0 10^4	22 23
W3(H_2O)	2.2	220	0.014	10	1.0 10^5	24
W51 e2[a]	8.0	140	0.093	200–400	1.5 10^6	25 26
W51 e8[a]	8.0	130	0.096	<200	''	25 26
W51 N-Dust	8.0	200	0.1	400	''	25 26

[a] with embedded UC HII region
[b] derived from $C^{18}O$

References: (1) Wright et al. 1996. (2) Wilson et al. 1996. (3) Peng et al. 1993. (4) Akeson and Carlstrom 1996. (5) Gómez et al. 1991. (6) Cesaroni et al. 1994a. (7) Hofner et al. 1996. (8) Cesaroni et al. 1998. (9) Olmi et al. 1996b. (10) Garay et al. 1993 (11) Ho et al. 1994. (12) Keto et al. 1988. (13) Cesaroni pers. comm. (14) Hofner et al. in prep. (15) Millar, Macdonald and Gibb 1997 (16) Hunter et al. 1998 (17) Garay et al. 1990 (18) Heaton et al. 1989 (19) Hunter et al. 1997. (20) Hofner et al. 1998. (21) Cesaroni et al. 1997. (22) Mangum et al. 1991. (23) Mangum et al. 1992. (24) Wyrowski et al. 1997. (25) Zhang and Ho 1997. (26) Zhang et al. 1998.